

Atmospheric neutrino observations and flavor changing interactions

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Flavor changing (FC) neutrino-matter interactions can account for the zenith-angle dependent deficit of atmospheric neutrinos observed in the SuperKamiokande experiment, without directly invoking neither neutrino mass, nor mixing. We find that FC ν_μ -matter interactions provide a good fit to the observed zenith angle distributions, comparable in quality to the neutrino oscillation hypothesis. The required FC interactions arise naturally in many attractive extensions of the Standard Model.

Neutrinos produced as decay products in hadronic showers from cosmic ray collisions with nuclei in the upper atmosphere [1] have been observed by several detectors [2–7]. Although the absolute fluxes of atmospheric neutrinos are largely uncertain, the expected ratio (μ/e) of the muon neutrino flux ($\nu_\mu + \bar{\nu}_\mu$) over the electron neutrino flux ($\nu_e + \bar{\nu}_e$) is robust, since it largely cancels out the uncertainties associated with the absolute flux. In fact, this ratio has been calculated [1] with an uncertainty of less than 5% over energies varying from 0.1 GeV to 100 GeV. In this resides our confidence on the long-standing atmospheric neutrino anomaly.

Although the first iron-calorimeter detectors in Fréjus [2] and NUSEX [3] reported a value of the double ratio, $R(\mu/e) = (\mu/e)_{\text{data}}/(\mu/e)_{\text{MC}}$, consistent with one, all the water Cherenkov detectors Kamiokande [4], IMB [5] and SuperKamiokande [6] have measured $R(\mu/e)$ significantly smaller than one. Moreover, not long ago, the Soudan-2 Collaboration, also using an iron-calorimeter, reported a small value of $R(\mu/e)$ [7], showing that the so-called atmospheric neutrino anomaly was not a feature of water Cherenkov detectors.

Recent SuperKamiokande high statistics observations [6] indicate that the deficit in the total ratio $R(\mu/e)$ is due to the number of neutrinos arriving in the detector at large zenith angles. Although e -like events do not present any compelling evidence of a zenith-angle depen-

dence, the μ -like event rates are substantially suppressed at large zenith angles.

The $\nu_\mu \rightarrow \nu_\tau$ [6,8] as well as the $\nu_\mu \rightarrow \nu_s$ [8,9] oscillation hypothesis provide an appealing explanation for this smaller-than-expected ratio, as they are simple and well-motivated theoretically. This led the SuperKamiokande Collaboration to conclude that their data provide good evidence for neutrino oscillations and neutrino masses.

In this letter we give an alternative explanation of the atmospheric neutrino data in terms of FC neutrino-matter interactions [10–14]. We show that even if neutrinos have vanishing masses and/or the vacuum mixing angle is negligible, FC neutrino matter interactions can still explain the SuperKamiokande data.

There are attractive theories beyond the SM where neutrinos are naturally massless [15] as a result of a protecting symmetry, such as B-L in the case of supersymmetric $SU(5)$ models [16] and the model proposed in [17] or chiral symmetry in theories with extended gauge structure such as $SU(3)_c \otimes SU(3)_L \otimes U(1)_N$ (331) models [18]. The simplest example of this mechanism was first noted in an $SU(2) \times U(1)$ model proposed in Ref. [12,17] where a singlet Dirac lepton at the TeV scale is added sequentially to the SM in such a way that neutrinos remain massless due to an imposed B-L symmetry. The flavor mixing amongst the massless neutrinos in the leptonic charged current can not be rotated away despite neutri-

nos being massless.

In general, models of neutrino mass are a natural source of FC neutrino-matter interactions. Seesaw-type models of neutrino mass have non-diagonal neutral current couplings of the Z to mass eigenstate neutrinos [19] that may lead to new FC neutrino-matter interactions. Models of radiative generation of neutrino mass [20] typically contain additional FC neutrino-matter interactions from scalar exchanges. An example of this class are the supersymmetric models with broken R -parity [21]. Models with extended gauge structure, such as $E(6)$ models [15], may also lead to FC neutrino-matter interactions.

Here we focus on a massless neutrino conversion scenario as an explanation of the atmospheric neutrino data based on FC neutrino-matter interactions which induce $\nu_\mu \rightarrow \nu_\tau$ transition. Our results can be extended to the massive neutrino case. However, we stress that the present atmospheric neutrino data does not necessarily provide evidence for neutrino mass. Moreover, the existence of attractive theories where FC neutrino-matter interactions do not imply neutrino mass, makes this possibility especially elegant. From a phenomenological point of view FC interactions of neutrinos can induce flavor transitions when neutrinos travel through matter [10] irrespective of neutrino mass. In both massive [22] and massless [12,13] cases conversions can be resonant, however the properties of the conversion are totally different. Massless neutrino conversions would be energy-independent and would affect atmospheric neutrinos as well as anti-neutrinos, converting $\nu_\mu \rightarrow \nu_\tau$ together with $\bar{\nu}_\mu \rightarrow \bar{\nu}_\tau$. Remarkably the present double ratio data does not show a significant energy dependence [6,23].

The presence of FC neutrino-matter interactions implies a non-trivial structure of the neutrino evolution Hamiltonian in matter. The evolution equations describing the $\nu_\mu \rightarrow \nu_\tau$ transitions in matter are given as [12,13]:

$$i \frac{d}{dr} \begin{pmatrix} \nu_\mu \\ \nu_\tau \end{pmatrix} = \sqrt{2} G_F \begin{pmatrix} 0 & \epsilon_\nu n_f(r) \\ \epsilon_\nu n_f(r) & \epsilon'_\nu n_f(r) \end{pmatrix} \begin{pmatrix} \nu_\mu \\ \nu_\tau \end{pmatrix}, \quad (1)$$

where, $\nu_a \equiv \nu_a(r)$, $a=\mu, \tau$ are the probability amplitudes to find these neutrinos at a distance r from their creation position, $\sqrt{2} G_F n_f(r) \epsilon_\nu$ is the $\nu_\mu + f \rightarrow \nu_\tau + f$ forward scattering amplitude and $\sqrt{2} G_F n_f(r) \epsilon'_\nu$ is the difference between the $\nu_\tau - f$ and $\nu_\mu - f$ elastic forward scattering amplitudes, with $n_f(r)$ being the number density of the fermions which induce such processes.

The use of the FC ν_e -matter interactions was previously suggested in connection with the solar neutrino problem [13,14]. Recently an attempt was made [24] to extend this idea in order to account also for the atmospheric neutrino data, but the fit obtained in this paper is not as good as our atmospheric neutrino fit or the solar fit in Ref. [14]. Moreover, some of the results in Table III of Ref. [24] seem inconsistent. On the other hand Ref. [25] includes exotic flavor-conserving ν_τ interactions

plus neutrino masses in order to account for the atmospheric and LSND data, but without a detailed fit. We have decided to postpone the detailed analysis (within the present scenario) of the LSND, as well as of the solar neutrino data, for a future work. Here we show that FC ν_μ -matter interactions can explain the atmospheric neutrino zenith angle anomaly, without introducing neutrino masses and/or mixing.

For our phenomenological approach let us simply assume the existence of a tree-level process $\nu_\alpha + f \rightarrow \nu_\beta + f$ with amplitude proportional to $g_{\alpha f} g_{\beta f} / 4m^2$, where α and β are flavor indices, f stands for the interacting elementary fermion (charged lepton, d -like or u -like quark) and $g_{\alpha f}$ is the coupling involved in the vertex where a ν_α interacts with f through a scalar or vector boson of mass m . The evolution equations which describe the $\nu_\mu \rightarrow \nu_\tau$ transitions in matter may be written generically as in Eq. (1), where:

$$\epsilon'_\nu = \frac{|g_{\tau f}|^2 - |g_{\mu f}|^2}{4m^2 \sqrt{2} G_F}, \quad \text{and} \quad \epsilon_\nu = \frac{g_{\tau f} \cdot g_{\mu f}}{4m^2 \sqrt{2} G_F}. \quad (2)$$

Note also that, in the absence of neutrino mass as we are assuming, anti-neutrino transitions $\bar{\nu}_\mu \rightarrow \bar{\nu}_\tau$ are governed by precisely the same evolution matrix in Eq. (1). We have calculated the transition probabilities of $\nu_\mu \rightarrow \nu_\tau$ ($\bar{\nu}_\mu \rightarrow \bar{\nu}_\tau$) $P(\epsilon_\nu, \epsilon'_\nu)$, ($P(\epsilon_{\bar{\nu}}, \epsilon'_{\bar{\nu}})$) as a function of the zenith angle by numerically solving the evolution equation using the density distribution in [26] and a realistic chemical composition with proton/neutron ratio 0.497 in the mantle and 0.468 in the core [27]. For the sake of simplicity we have assumed that $\epsilon \equiv \epsilon_{\bar{\nu}} = \epsilon_\nu$ and $\epsilon' \equiv \epsilon'_{\bar{\nu}} = \epsilon'_\nu$, so there are only two free parameters in the analysis. We have used these probabilities to compute the theoretically expected number of μ - and e -like events (N_μ and N_e) as a function of the two parameters, ϵ and ϵ' , for each of the 5 zenith angle bins both for the sub-GeV and for the multi-GeV SuperKamiokande data. Following Refs. [8,28], we fit separately the μ - and e -like events taking into account the correlation of errors. The calculated numbers of events, N_μ and N_e , as functions of zenith angle, have been compared with the 535-days SuperKamiokande data sample in order to determine the allowed regions of ϵ and ϵ' from a χ^2 fit. We set our normalization assuming that the relevant neutrino interaction in the Earth is only with down-type quarks. Any other scenario can be obtained from our results by rescaling the ϵ -parameters.

In Fig.(1) we present a contour plot of the regions allowed by the SuperKamiokande data. The contour plots (a), (b) and (c) correspond to the regions allowed by the sub-GeV, multi-GeV and combined SuperKamiokande data, respectively. These contours are determined by the conditions $\chi^2 = \chi_{min}^2 + \Delta\chi^2$ where $\Delta\chi^2 = 4.6, 6.0, 9.2$ for 90, 95 and 99 % C. L., respectively.

In the parameter region we have considered, i.e., ϵ and

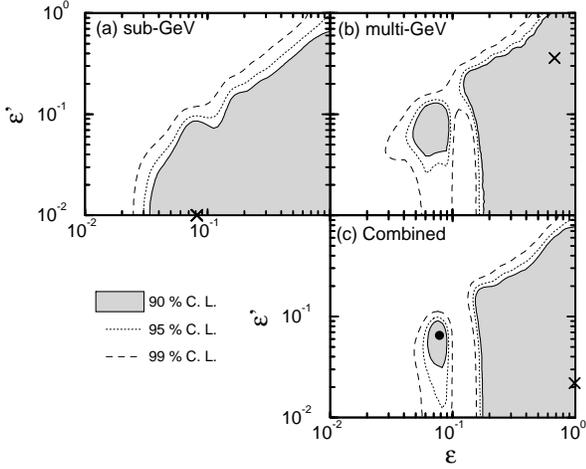


FIG. 1. Allowed region for ϵ and ϵ' for SuperKamiokande (a) sub-GeV (b) multi-GeV and (c) combined events in the massless-neutrino scenario. The best fit points for each case is indicated by the crosses.

ϵ' in the interval $[0.01, 1.0]$, we found that $\chi^2_{min} = 6.3$ and 6.4 for the sub-GeV and multi-GeV samples (8 d. o. f. corresponding to 10 data points minus two free parameters). These minima are obtained for $(\epsilon, \epsilon') = (0.08, 0.01)$ and $(0.68, 0.36)$, respectively, as indicated by the crosses in Fig.(1). For the combined case, $\chi^2_{min} = 14.7$ (18 d.o.f) for $(\epsilon, \epsilon') = (0.99, 0.02)$. In the combined case the local best fit point ($\chi^2 = 16.9$ for $(\epsilon, \epsilon') = (0.08, 0.07)$) in the “island” determined by the 90 % C. L. curve is also indicated by a filled circle. This point is interesting because it still gives a good fit to the data with a relatively small value for the FC parameter ϵ . We find also that the χ^2 is relatively flat along the ϵ' axis around the best fit point. The allowed regions can be qualitatively understood in the approximation of constant matter density. The conversion probability in this case is:

$$P(\nu_\mu \rightarrow \nu_\tau) = \frac{4\epsilon^2}{4\epsilon^2 + \epsilon'^2} \sin^2\left(\frac{1}{2}\eta L\right), \quad (3)$$

where $\eta = \sqrt{4\epsilon^2 + \epsilon'^2} \sqrt{2G_F n_f}$. For $n_f = n_d \approx 3n_e$ and $\epsilon' < \epsilon$, the oscillation length in matter is given by:

$$L_{osc} = \frac{2\pi}{\eta} \approx 1.2 \times 10^3 \left[\frac{2 \text{ mol/cc}}{n_e} \right] \left[\frac{1}{\epsilon} \right] \text{ km}. \quad (4)$$

From Eq. (3) one can see that in order to have a large transition probability one must be in the region $\epsilon' \lesssim \epsilon$ and $\eta \lesssim \pi/R_\oplus$. This last condition leads to a lower bound on ϵ . The island in Fig. 1.(b) corresponds to $\eta \sim \pi/R_\oplus$.

In Fig.(2) we give the expected zenith angle distribution of μ -like sub-GeV events (a) and multi-GeV events (b) evaluated with our Monte Carlo program for the best fit points determined above. Our results clearly indicate an excellent fit for the μ -like events showing that they are highly depleted at $\cos\theta = -1$ with respect to the SM

prediction. Note that, except for the assumption that the FC ν_μ -matter interaction involves d -quarks, our result is quite general, since we have not explicitly considered any particular model as the origin of the FC neutrino-matter interaction. Note that e -like events are not affected by $\nu_\mu \rightarrow \nu_\tau$ transition.

What can we say about the required strength of the neutrino-matter interaction in order to obtain a good fit of the observed data? From our results and Eq. (2) we see that for masses $m \approx 200$ GeV we need at least $g_{\tau f} \cdot g_{\mu f} \sim 0.1$ for the the mixing term ϵ . Similarly our best fit ϵ' value implies $|g_{\tau f}|^2 - |g_{\mu f}|^2 \sim 0.1$. While these values are relatively large, they are both weak-strength couplings. Moreover they are consistent with present experimental bounds, for example from universality of the weak interaction which is manifestly violated by Eq. (1).

For the purpose of illustrating this explicitly let us consider for the moment the supersymmetric model with broken R -parity [21] as a way to parameterize the FC neutrino-matter interaction. In this case the FC ν_μ -matter interactions are mediated by a scalar down-type quarks, \tilde{d}_j , so that we need only to check the couplings where a d -quark and a μ - or τ -neutrino is involved, *i.e.* $g_{id} \approx \lambda'_{ij1}$, $i = 2, 3$. λ'_{ijk} are the coupling constants in the broken R -parity superpotential $\lambda'_{ijk} L_i Q_j D_k^c$, where L , Q and D are standard superfields, and $4\sqrt{2}G_F\epsilon = \left| \sum_j \lambda'_{3j1} \lambda'_{2j1} / \tilde{m}_{\tilde{d}_j}^2 \right|$.

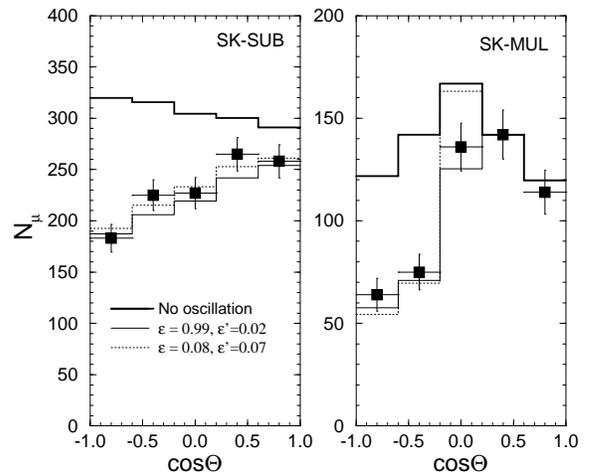


FIG. 2. Best-fit zenith angle distributions in the massless-neutrino FC scenario (thin-solid, dotted lines) versus no-oscillation hypothesis (thick-solid line). The SuperKamiokande data are indicated by the crosses.

Constraints on the magnitude of such FC interactions in broken R -parity models have been given in [29]. The precision tests imply that individually these couplings are not severely constrained. The most stringent limit to the values of the relevant FC quantities comes from limits on the FC tau decay $\text{BR}(\tau^- \rightarrow \rho^0 + \mu^-) < 6.3 \times 10^{-6}$ [30]

which implies that $|\sum_j \lambda'_{3j1} \lambda'_{2j1} (100 \text{ GeV}/\tilde{m}_{\tilde{u}_j})^2| < 3.1 \times 10^{-3}$. Although a certain degree of fine-tuning is needed in order to verify this constraint we find that the required strength of FC ν_μ -matter interaction is consistent with all present data. One should note that broken R -parity supersymmetric models typically lead to neutrino masses which could be large. However one may suppress them via fine-tuning or some additional symmetry.

In summary, we have demonstrated that flavor changing ν_μ -matter interactions can account for the zenith-dependent deficit of atmospheric neutrinos observed in the SuperKamiokande experiment, without directly invoking neutrino masses and mixing. It provides a fit of the observations which is significantly better than the no-oscillation hypothesis and of similar quality as the usual $\nu_\mu \rightarrow \nu_\tau$ oscillation hypothesis. The required FC interaction can arise in many attractive extensions of the SM and is consistent with all present constraints.

The above FC mechanism can also be tested at future Long Baseline experiments. From Eq. (3), using $n_e \sim 2 \text{ mol/cc}$, $\epsilon \sim 1$ (0.1) and $\epsilon' < \epsilon$, for the planned K2K experiment [31] one gets $P(\nu_\mu \rightarrow \nu_\tau) \sim 0.4$ (0.004) while for MINOS [32] one finds $P(\nu_\mu \rightarrow \nu_\tau) \sim 0.9$ (0.2).

The existence of a massless neutrino explanation of the atmospheric neutrino anomaly may play an important theoretical role in model-building, especially if one wants to account for all other hints for non-standard neutrino properties, namely the solar neutrino data, the LSND result, and the possible role of neutrinos as dark matter in the Universe [15].

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